

# Multi-planetary Systems from Simulated TESS Transit Timing Variations

*Lucas Hellström*

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Lund Observatory  
Lund University



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Supervisor: Alexander Mustill

Lund Observatory  
Box 43  
SE-221 00 Lund  
Sweden

## Abstract

A transit is a phenomenon where a planet passes between its host star and an observer blocking out part of the light from the star. This decrease can be measured and used to gain information about the planet. The Kepler and TESS telescopes are examples of space telescopes using the transit method to detect exoplanets. For systems with more than one planet around the same star variations between the time a planet takes to transit might occur. These variations, called Transit Timing Variations, or TTVs, can be used to gain information about additional planets in the system.

TESS recently launched and this paper uses data from Sullivan et al. (2015) and Kepler data from Q1-Q17 DR25 from the NASA Exoplanet Archive (Thompson et al. 2018) to create artificial systems like those that TESS might observe to predict what kind of results it might find. These systems are simulated using TTVFast (Deck et al. 2014) to obtain TTV signals which are then used to create a sky map to show the fraction of systems showing TTV signals in a given sample.

TESS have coverage of the ecliptic poles for a whole year which results in that many systems showing considerable TTV signals being located at the poles although there are systems located outside the poles still showing TTV signals. In the range of declination from about  $-40^\circ$  to  $40^\circ$  CHEOPS will be able to further study the objects. As CHEOPS will study already known objects predictions for the results from TESS can be used to estimate the number of targets CHEOPS might study, how much observation time is required for each object and how to prioritise the most interesting systems. Using information about the planetary radius, period and stellar mass obtained from short-term observations the long term TTV amplitude can be estimated. From the results of this paper it is expected that CHEOPS can find a TTV signal from about every fourth transiting planet and the probability for a CHEOPS detectable TTV signal is high where the period ratio is below 2 but low for a higher fraction.





## Populärvetenskaplig beskrivning

När vi letar efter exoplaneter finns det ett antal olika metoder för att hitta dem. Den mest framgångsrika är transitmetoden där ljusstyrkan hos en stjärna studeras under en längre tid. När en planet passerar mellan sin stjärna och en observatör kan en minskning i stjärnans ljusstyrka ses. Uppreras detta i regelbunda intervall kan slutsatsen att det finns en planet runt stjärnan dras. Genom att studera minskningen i ljusstyrka kan storleken på planeten beräknas vilket kombinerat med massan som fås av andra metoder ge en insikt i hur och vad planeten är uppbyggd av. En transit är detta fenomen då en planet passerar mellan stjärnan och en observatör.

Genom att jämföra tiden mellan varje transit för en planet kan ibland variationer ses, vilket kallas Transit Timing Variations eller förkortat TTV. Detta beror på att det finns fler planeter runt stjärnan som med hjälp av gravitationskraften accelererar eller decelererar planeten som bevakas. Detta resulterar i att det är möjligt att hitta planeter som genom andra metoder är osynliga.

Keplerteleskopet är ett rymdbaserat teleskop som använder transitmetoden för att hitta exoplaneter. Det har sedan 2009 hittat över 1000 bekräftade exoplaneter vilket gör den till det hittills mest framgångsfulla uppdraget i jakten på exoplaneter. TESS, vilket står för Transiting-Exoplanet Survey Satellite, är ett teleskop som sköts upp den 18de april 2018 och använder transitmetoden för att hitta exoplaneter. TESS kommer bli det första rymdbaserade teleskopet att studera hela himlen och kommer observera över 200 000 stjärnor under uppdragets ursprungliga längd på två år.

Detta projekt kommer använda data från Keplerteleskopet för att simulera data från TESS för att sedan använda den datan för att leta efter TTV signaler. Detta ska ge en uppfattning om hur många system som har fler än en planet inom ett givet område på himlen.



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# Chapter 1

## Introduction

When observing stars in the search for exoplanets a few different methods can be used. The most successful method so far is the transit method which measures the brightness of a star for a long period. If a planet passes between the star and the observer it will block out part of the light and the brightness will decrease. If these decreases occur at regular intervals the conclusion that the reason for this is an exoplanet can be drawn. By studying the amount of light the planet blocks out the radius of the planet can be found; combined with the mass of the planet obtained from different methods the approximate density of the planet can be calculated. This gives information about the structure and composition of the planet.

In a system with multiple planets around the same star, the planets will affect each other through their gravitational pull. This results in the planets accelerating or decelerating depending on the relative positions of the planets. As this is happening the time of one orbit may differ and by studying these variations, planets which may not be possible to detect through the transit method can be detected (Nesvorný et al. 2012). TTVs can also be used to calculate the mass of the transiting planet and are commonly used to confirm exoplanet candidates as real planets.

This paper will simulate data from the TESS telescope to search for these transit timing variations in order to determine the approximate fraction of multi-planetary systems in a given sample. The amplitude dependence on observation time is plotted and compared to an analytically approximation to understand the importance of longer observation times in order to detect TTVs. Finally systems in the range where CHEOPS are able to observe will further be studied in order to determine the fraction of detectable TTV signals based on the radius and period ratio of the planets in the system which will predict what kind of systems are best suited for follow-up observations.

## 1.1 Transits

A planet in orbit around its host star may sometimes cross the line of sight of an observer. When this happens a slight decrease in the star's brightness can be measured. This is called a transit and is today used as a main method to discover exoplanets. From transits the radius of the planet can be determined but it can also be used to find additional planets around the host star which may not be transiting. This will be discussed in section 1.1.1. With the radius known from the transit method and the mass obtained from for example the TTV signal or different methods such as, for example the radial velocity method, the density of the planet can be calculated. The density is important to understand what the planet is made of and the structure of it.

### 1.1.1 Variations

For a system with a single planet around a star the period of the planet is more or less perfectly periodic with no visible variations but, when measuring the time between a planet's transits one may discover variations in the period which are called Transit-Timing Variations or TTVs. These variations arise from another planet in the system whose gravitational pull accelerates or decelerates the observed planet which results in increased or decreased transit times. The amplitude of the TTV signal is dependent on the mass of the planets and the ratio of the periods between them, some ratios give rise to a resonance which increases the amplitude of the TTV (Nesvorný & Morbidelli 2008). TTV signals are typically sinusoidal (Lithwick et al. 2012) with a period called the super-period. This super-period is heavily dependent on period fraction of adjacent planets and the mass of planets in the system. The super-periods are often in the range of 10 to 100 days, which means that the period of many TTV signals are too long to be observed by a single TESS observation period and requires further measurement. An advantage of studying transits in search for TTVs is that planets which do not transit their star can be discovered through TTVs (Nesvorný et al. 2013). As most planets do not transit their star this can increase the number of known exoplanets drastically.

## 1.2 Kepler

The Kepler satellite launched in spring 2009 on a mission to study stars in a small patch in the sky to discover Earth-sized exoplanets within the habitable zone, where

liquid water can exist on the planetary surface. The brightness of a large number of stars are measured and then analyzed in order to detect transiting exoplanets.

Kepler started by looking at a very small patch of the sky but in July 2012 one of the four wheels used to keep the patch in focus broke. The telescope requires at least three wheels to function which kept the mission alive. In May 2013 a third wheel failed which resulted in the telescope no longer being able to collect data. The satellite was nonfunctional until the so-called "Second Light (K2)" in early 2014. This mission would use the telescopes remaining two wheels to study stars over a much larger area but for shorter periods. (Borucki 2017)

Kepler have found over 4500 exoplanet candidates (Borucki 2017) to this day. This amount makes Kepler the most successful exoplanet hunting mission to this date.

## 1.3 TESS

The Transiting Exoplanet Survey Satellite, TESS, is a satellite which was launched on the 18th of April 2018. The satellite is equipped with four cameras which will study the brightness of over 200 000 stars over a two year period. It is the first all-sky transit survey taking place in space. (Ricker et al. 2014)

TESS will study the whole sky by splitting it into 26 sectors, 13 in the southern hemisphere and 13 in the northern hemisphere, which are observed for 27 days each. An illustration of this can be seen in figure 1.1 where the number of times TESS will observe each sector are shown. TESS will observe the southern hemisphere during the mission's first year, it will then rotate and observe the northern hemisphere for another year.

## 1.4 CHEOPS

The CHaracterising ExOPlanets Satellite, CHEOPS, is a satellite currently in development which is planned to launch in early 2019. It will study already known bright stars hosting planets for transits with ultrahigh precision photometry (Broeg et al. 2013).

TESS will discover several hundred systems and give a fair estimation of the planetary radius and mass of transiting planets in these systems. TESS will also be able to find TTV signals for systems with a low super-period. CHEOPS will be able to use this data to look at planets with mass between  $1 M_{\oplus}$  and  $20 M_{\oplus}$  in order to obtain more accurate radii measurements which in turn can be used to determine the



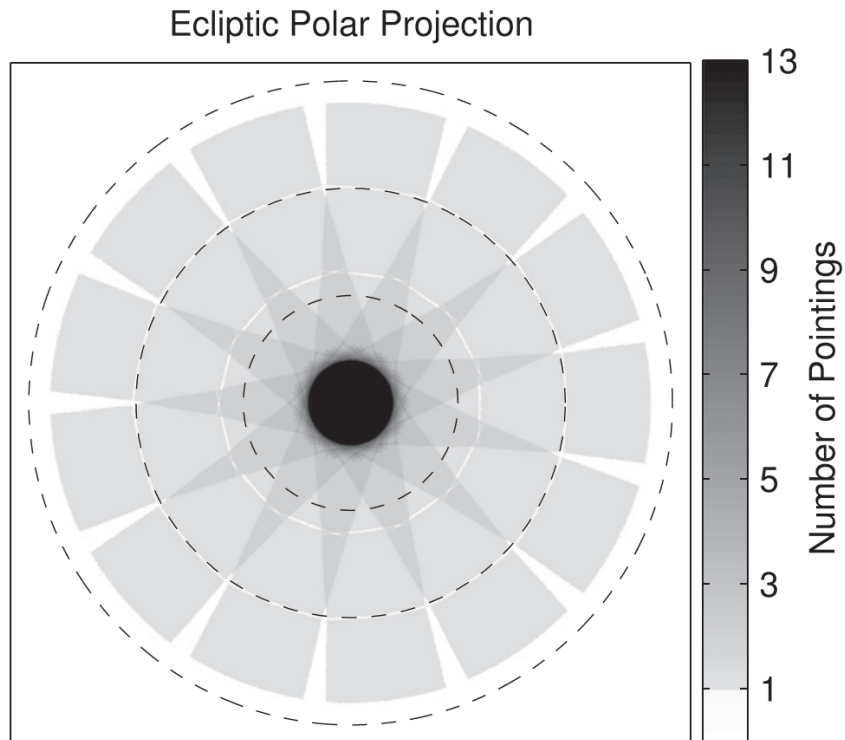


Figure 1.1: Illustration of the number of times TESS will observe each sector in the sky.

Source: Sullivan et al. (2015), figure 1

structure which might help in understanding the formation and evolution of planets in this mass range. CHEOPS will also be able to make follow-up TTV measurements for more accurate TTV data.

Due to limitations on CHEOPS it is not able to study the whole sky and is limited to a range of about 40 degrees above and below the ecliptic. This creates a problem with candidates obtained from TESS as the systems in this range are only observed for one or two observation periods while the ecliptic poles have all year coverage.

# Chapter 2

## Method

This paper uses results from Sullivan et al. (2015) in which the properties of the transiting planets that TESS will detect are predicted. The results from Sullivan et al. (2015) are based on Monte Carlo simulations, occurrence rates of planets obtained from Kepler and models of the performance of the TESS cameras. A limitation of the simulated detections by Sullivan et al. is that, although they report the number of planets per star, parameters for only one planet are actually provided. However, we can use the large number of multi-planet systems discovered by Kepler to create artificial systems based on the results from Sullivan et al. and the Kepler dataset. These systems are then simulated to obtain TTVs. The steps taken are listed below:

1. For each planet in the Sullivan et al. catalogue a similar planet within 10% in radius and period are found in the Kepler dataset.
2. The systems are then filled with additional planets taken from the Kepler archive and modified based on the ratio of the radius and period of the first Sullivan and Kepler planet. The planetary period and signal-to-noise ratio for all planets are checked in order to ensure that they are detectable by TESS.
3. These planets are then assigned a mass depending on the radius of the planet (Sullivan et al. 2015) and an eccentricity based on a Rayleigh distribution as none of those parameters are obtained from the light curves.
4. The systems are then simulated using TTVFast (Deck et al. 2014) and the resulting transit times are analyzed. First a linear fit needs to be applied and subtracted to remove the period of the planet and thus obtain the variation. The amplitude of the variation is obtained from average of the highest and lowest value.

5. The same procedure, mass-radius relation and distribution are applied to Kepler systems and the results compared to the TTV catalogue of Ofir et al. (2018) to verify that the methods and assumptions give realistic results.
6. The error of the results are calculated from Holman & Murray (2005) and is dependent on the magnitude of the host star, the ratio of radius between the planet and host star and the transit duration.
7. To ensure that the systems are realistic they are simulated for stability over a long time. This is done using WHFast (Rein & Tamayo 2015) and IAS15 (Rein & Spiegel 2015).
8. The amplitude dependence on observation time are plotted and compared to an analytical approximation, giving an idea of how the amplitude increases over time. This shows the importance of longer observation times in order to obtain full TTV signals.
9. Systems in the CHEOPS observation range,  $-40 < \beta < 40$ , are further studied in order to obtain the fraction of planets showing a TESS or CHEOPS detectable TTV signal for a given radius and period ratio.

## 2.1 Simulation of TESS objects

Sullivan et al. (2015) provides one planet per system. By combining data from the Kepler dataset<sup>1</sup> and the results from Sullivan et al. (2015), we simulate artificial TESS systems to obtain TTV signals. Sullivan et al. (2015) uses the occurrence rates of planets around a host star with effective temperature  $T_{eff} < 4000$  K from the results reported by Dressing & Charbonneau (2015) while for  $T_{eff} > 4000$  K the occurrence rates are obtained from Fressin et al. (2013). Because of this the planets are separated into two groups based on the effective temperature of the host star. One group with effective temperature below 4000 K and one for higher.

For each planet in the Sullivan et al. (2015). catalogue, a similar planet, in radius and period, are selected from the Kepler dataset. The planets in this system are then verified that their periods are not longer than the observation time of TESS. To further ensure detectability the signal-to-noise ratio, SNR, is calculated:

$$SNR = \frac{T_{depth}}{N_6} \quad (2.1)$$

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<sup>1</sup><https://exoplanetarchive.ipac.caltech.edu/index.html>

where  $T_{depth}$  is the transit depth and  $N_6$  is the noise from 6 hours of observations. The noise is based on photon counting which follow Poisson statistics. The noise is thus the square root of the photon count. As a baseline of 6 hours is used the noise is  $\sqrt{6\Gamma}$ :

$$SNR = \left(\frac{R_p}{R_\star}\right)^2 \times \sqrt{6\Gamma} \quad (2.2)$$

where  $R_p$  and  $R_\star$  are the radius of the planet and host star and  $\Gamma$  is the photon count per hour. From Sullivan et al. (2015) the lowest SNR where a signal is detectable securely without introducing false positives over the whole mission is about 7.3, this is used as the lower limit for the SNR in this paper.

The ratio of radius and period between the two planets are calculated and multiplied with the Sullivan planets radius and period. This results in that the two planets are identical in radius and period. These ratios are then applied to the rest of the planets in this selected Kepler system to create an artificial system of planets.

The mass of the planets are required to simulate the orbits and TTV signals of the planets and are approximated using equations 2.3 and 2.4 obtained from Weiss et al. (2013):

$$M_p = M_\oplus \left[ 0.440 \left(\frac{R_p}{R_\oplus}\right)^3 + 0.614 \left(\frac{R_p}{R_\oplus}\right)^4 \right] \quad (2.3)$$

for planets with  $R_p < 1.5 R_\oplus$  where  $R_\oplus$  is the radius of Earth and  $M_\oplus$  is the mass of Earth. For planets with  $R_p \geq 1.5$  the equation changes to:

$$M_p = 2.69M_\oplus \left(\frac{R_p}{R_\oplus}\right)^{0.93} \quad (2.4)$$

When setting up the system the mean anomaly is required. This specifies the positions of all planet in a given system at a snapshot in time and is used as a starting point for the simulation. It is acquired from the number of transits and the orbital period of the planet by using a reference point specified when a planet is directly in front of the host star as seen from an observers point of view. This is illustrated in figure 2.1 for a two planet system. This reference point corresponds to a mean anomaly of  $90^\circ - \omega$ , where  $\omega$  is the argument of periapsis, at some time  $T_i$ , where  $i$  is the number of the planet in the system. For a two planet system this means that  $M_1 = 90^\circ - \omega$  at some time  $T_1$  and  $M_2 = 90^\circ - \omega$  at some time  $T_2$ . In order to calculate the mean anomaly at some time a time reference point is defined as  $t = 0$  and the goal is the calculate the mean anomaly of some planet  $i$ . This can be done by using the mean anomaly at  $t = T_i$  and subtracting the number of degrees,  $\xi$ , the

planet have traveled since then:

$$M_i(t = 0) = M_i(t = T_i) - \xi \quad (2.5)$$

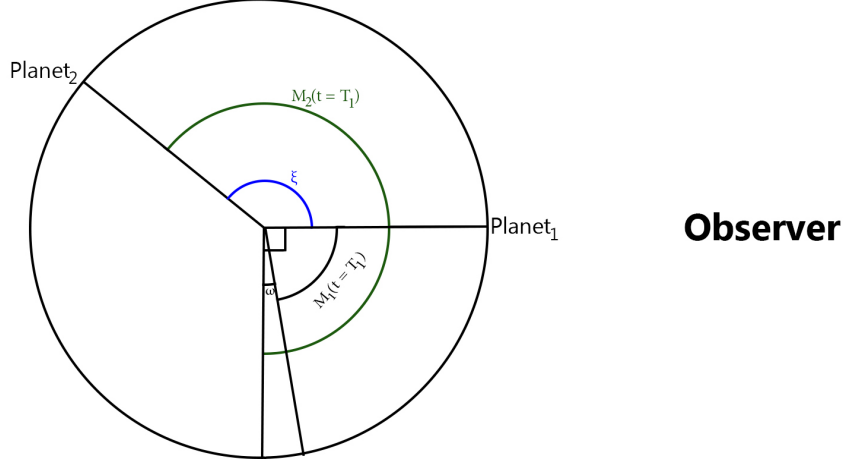


Figure 2.1: Illustration of the mean anomaly of the second planet at  $t = T_1$  when one of the planets is transiting.

The mean anomaly at time  $T_i$  is  $M_i(t = T_i) = 90^\circ - \omega$  and the number of degrees traveled is  $360^\circ$  multiplied by the number of orbits since  $t_0$ :

$$M_i(t = 0) = 90 - 360 \frac{T_{epoch}}{P_i} - \omega \quad (2.6)$$

where  $T_{epoch}$  is the transit epoch and  $P_i$  is the period of the planet. The argument of periapsis is obtained from a uniform distribution where  $0 < \omega < 360$ .

TTVFast also requires inclination, eccentricity and longitude of the ascending node. These cannot be simply obtained and need to be assumed. For simplicity the inclination is assumed to be  $90^\circ$  for all planets which results in the longitude of the ascending node to be undefined. The eccentricity is obtained from a Rayleigh distribution (Van Eylen et al. 2018; Wu & Lithwick 2013) with mode  $\sigma = 0.03$  in order to not get eccentricities above 0.1 as that leads to unstable systems.

How long a system will be observed depends on where in the sky it is located. This time can be obtained from the Web TESS Target tool <sup>2</sup> if the right ascension,

<sup>2</sup><https://heasarc.gsfc.nasa.gov/cgi-bin/tess/webtess/wtm.py>

RA, and declination, dec, are known. This tool accepts csv files with RA and dec of the systems and outputs the number of times TESS will observe that system. Currently it only works with systems in the southern hemisphere i.e with ecliptic latitude below 0. Because of this our systems in the northern hemisphere have their latitude flipped and are such placed in the souther hemisphere when the data is uploaded to the TESS tool. The latitude is then flipped back before the data is used.

## 2.2 TTVFast

TTVFast is a program created by Deck et al. (2014) which simulates planetary systems using an n-body integrator. It requires information about the system in the form of:

- Gravitational constant in  $AU^3\text{day}^{-2}M_{\odot}^{-1}$
- Mass of the star [ $M_{\odot}$ ]

And also for each planet in the system:

- Period [days]
- Eccentricity
- Inclination [ $^{\circ}$ ]
- Longitude of ascending node [ $^{\circ}$ ]
- Argument of periapsis [ $^{\circ}$ ]
- Mean anomaly at the reference time [ $^{\circ}$ ]

Where longitude of ascending node and argument of periapsis are orbital elements. The reference time is the time of the start of the integration, in this paper:  $t_{ref} = 0$ . The program also requires parameters regarding the integration which are given in a setup file:

- Path to file containing info regarding the planets in the system.
- Reference time
- Time step which is 1/20 of the period
- Final time which in this paper is the duration of the integration
- Number of planets
- Input flag which specifies in which coordinate system the input parameters are given. This paper uses Jacobi coordinates which relate to a input flag = 0.

With all these known the system can be simulated and the output are given as a number of times when a transit occurred and the number of the transiting planet.

## 2.3 Simulation of Ofir objects

A paper written by Ofir et al. (2018) used observed objects from the Kepler dataset to obtain TTV signals. In order to determine the precision of the methods used in this paper, the systems observed by Ofir et al. (2018) are simulated using the same methods as in this paper for 4 years. The resulting amplitudes of the TTV signals are compared to those reported by Ofir et al.

## 2.4 Analyzing results from TTVFast

When all systems have been simulated with TTVFast the data needs to be analyzed to find the TTV signals. Assuming a constant period, the linear ephemeris needs to be calculated and subtracted from the transit times. This is done by making a linear fit to the times and subtracting this fit. In order to clearly see the amplitude of the TTV signals the times are corrected by subtracting the average time from every value which moves the middle of the graph to  $y = 0$ . In order to obtain the amplitude of the TTVs the average of the maximum and minimum transit time are calculated. This is used as the amplitude and the corrected transit times are plotted.

The position on the sky of the objects are of interest and are obtained from the Sullivan et al. catalogue. The positions are given in RA and dec and plotted in a sky map. These coordinates are given in the equatorial frame and converted to the ecliptic frame. When plotted, this produces sky maps which are colour coded according to the TTV amplitude and the multiplicity, i.e. the number of planets, of the systems.

## 2.5 Error estimation

TTV signals are only detectable and of interest if the amplitude is bigger than the error on the transit time. The error estimation on the transit time used in this paper comes from equation 3 in Holman & Murray (2005):

$$\sigma_t \approx \left[ (\Gamma t_T)^{-1/2} \left( \frac{R_p}{R_\star} \right)^{-3/2} \right] t_T \quad (2.7)$$

where  $\Gamma$  is the photon count rate of the observed star,  $t_T$  is the transit duration,  $R_p$  and  $R_\star$  is the radius of the planet and star in solar radii. For Kepler the photon count is  $\Gamma = 7.8 \times 10^8 10^{-0.4(V-12)} \text{ hr}^{-1}$  where  $V$  is the apparent magnitude of the

star. From Sullivan et al. (2015) the photon flux for TESS at magnitude  $I_c = 0$  is  $\Phi \approx 1.4 \times 10^6 \text{ s}^{-1} \text{ cm}^{-2}$  and  $\Phi \approx 1.4 \times 10^6 10^{-0.4I_c} \text{ s}^{-1} \text{ cm}^{-2}$  when  $I_c \neq 0$ . The diameter of a camera on TESS is 100 mm which gives an area of  $A_{\text{camera}} = \pi(50 \text{ mm})^2 = 7853.98 \text{ mm}^2 = 78.54 \text{ cm}^2$ . The photon count is the photon flux multiplied with the area of the camera:

$$\begin{aligned} \Gamma &= \Phi A_{\text{camera}} = 1.4 \times 10^6 10^{-0.4I_c} \text{ s}^{-1} \text{ cm}^{-2} \times 78.54 \text{ cm}^2 = \\ &1.10 \times 10^8 10^{-0.4I_c} \text{ s}^{-1} = 3.96 \times 10^{11} 10^{-0.4I_c} \text{ hr}^{-1} \end{aligned} \quad (2.8)$$

## 2.6 Stability simulations

The artificial systems that are created need to be checked for stability to determine if the assumptions and method of creating them results in realistic and stable systems. For this Rebound (Rein & Liu 2012), more specifically the WHFast integrator (Rein & Tamayo 2015), is used. WHFast uses a Wisdom-Holman symplectic integrator for long duration planetary system simulations up to order eleven. For some systems, close-encounters occur, due to the symplectic integrator and fixed timestep WHFast does not produce accurate results for close-encounters, therefore IAS15 (Rein & Spiegel 2015). is used to simulate these systems. IAS15 is a 15th-order integrator based on a Gauß-Radau quadrature with an adaptive timestep. This makes IAS15 much more accurate for close-encounters and high-eccentricity orbits.

WHFast requires the semi-major axis of each planet in each system. This is obtained from Kepler's law of periods:

$$a^3 = \frac{T^2 GM_\star}{4\pi^2} \Rightarrow a = \sqrt[3]{\frac{T^2 GM_\star}{4\pi^2}} \quad (2.9)$$

where  $a$  is the semi-major axis in AU and  $T$  is the period in years.

As the stability simulations take a long time not all systems can be checked for stability. The first systems are selected by their multiplicity in order to see that at least one system of each multiplicity is stable. The Hill radius is then calculated for each system:

$$r_{\text{Hill}} = \frac{a_1 + a_2}{2} \left( \frac{M_1 + M_2}{3M_\star} \right)^{1/3} \quad (2.10)$$

where  $a_{1,2}$  is the semi-major axis of the two inner most planets,  $M_{1,2}$  is the mass of the two planets and  $M_\star$  is the mass of the host star. This radius is then used to calculate the separation:

$$\Delta = \frac{a_2 - a_1}{r_{\text{Hill}}} \quad (2.11)$$



The systems with lowest separation are the systems most probable to be unstable (Chambers et al. 1996) and thus they are simulated for  $10^5$  to  $10^6$  years depending on the system, to ensure stability.

## 2.7 Amplitude dependence on time

For each transit of a planet in a system, an amplitude is calculated using the same method as in section 2.4. This is compared to an analytical approximation of the amplitude. The TTV signal are assumed to be a sine curve and the amplitude can therefore be approximated as:

$$A(t) = \left[ \frac{A_{max}}{2} \sin\left(\frac{2\pi t}{P}\right) - \frac{A_{max}}{2} \sin\left(\frac{2\pi t_{obs}}{P}\right) \frac{t}{t_{obs}} \right] \quad (2.12)$$

where  $A$  is the amplitude at time  $t$ ,  $A_{max}$  is the maximum amplitude of the curve,  $P$  is the period of the curve and  $t_{obs}$  is the total observation time of the selected system.

## 2.8 Objects in CHEOPS observation range

Systems in the are of the sky where CHEOPS can observe, called the observation range, are of interest for the upcoming TESS satellite. More specifically, being able to predict if a system shows TTV signals based on information obtainable on a short timescale is of interest. To obtain this, multiple artificial systems are created for each system in the CHEOPS range. For each system in this range, 100 additional systems are created with fixed planetary masses, periods and star mass but with varying eccentricities, arguments of periapsis and mean anomalies. These systems are then simulated using TTVFast for 1 year and analysed in the same way as in section 2.4.

The amplitude of each planet is plotted against the period fraction of the two closest planets. As this is done the ratio of periods of the planets in each system are calculated with respect to the two closest planets and then plotted against the radius of the planet. This plot is color coded with respect to the amplitude.

In order to obtain the fraction of systems with a detectable TTV signal for a given period fraction and planet radius a 2d histogram is created with these parameters. Where the color of each square in the grid corresponds to the probability of finding a detectable TTV signal obtained from the simulated TESS objects in the CHEOPS range.

# Chapter 3

## Results

### 3.1 Simulated TESS objects

The planetary radius of the artificial systems are plotted against the orbital periods in figure 3.1 where the multiplicity of the system are represented with colour and shape. Most systems consists of two or three planets and systems containing five or six planets are few. Almost all planets are larger than earth and have a far shorter period. Larger planets are easier to discover as they block out more of the light from the star and short period planets allows for more opportunities to observe the transit which results in that most discovered planets are large with short periods. The figure shows several small clumps of planets. This is due to that the same Kepler system can be sampled multiple times.

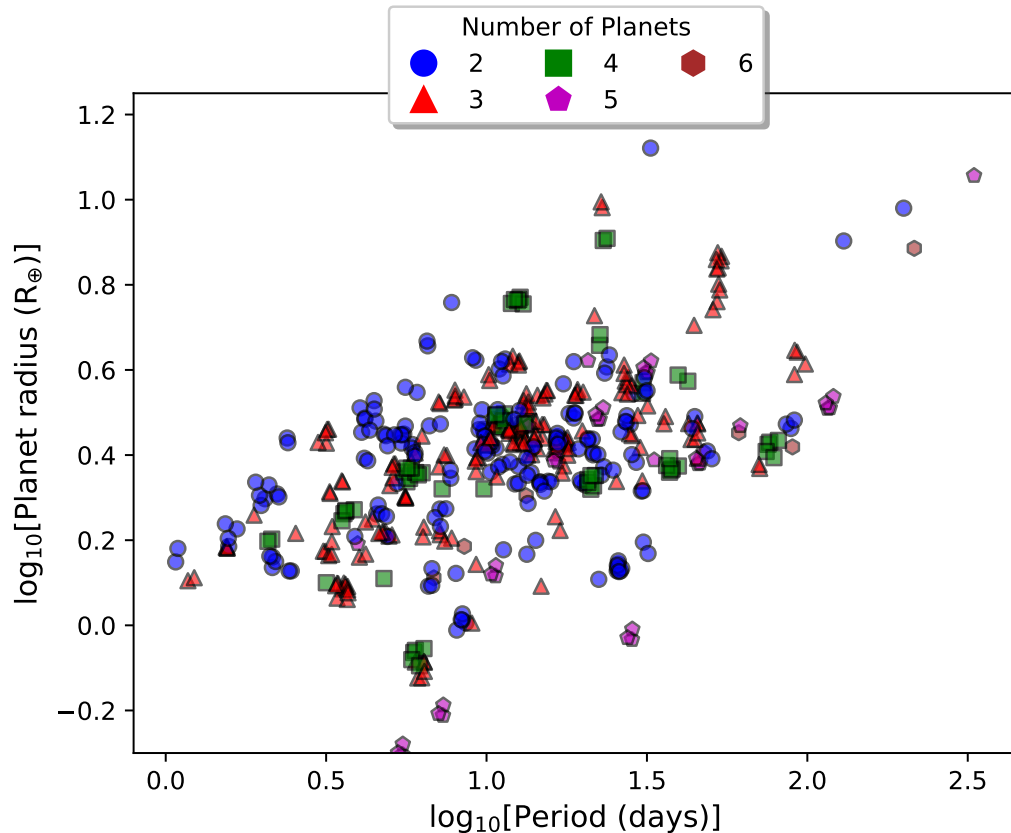


Figure 3.1: Radius distribution as a function of period for the simulated TESS objects where the colour and shape corresponds to the multiplicity of the system.

Figure 3.2 shows a sky map in the ecliptic frame of the simulated TESS objects and the number of times they are observed. The green area shows the approximate range where CHEOPS is able to observe the objects. As expected, the systems in the ecliptic poles are observed the most times and most objects in the CHEOPS range are only observed once.

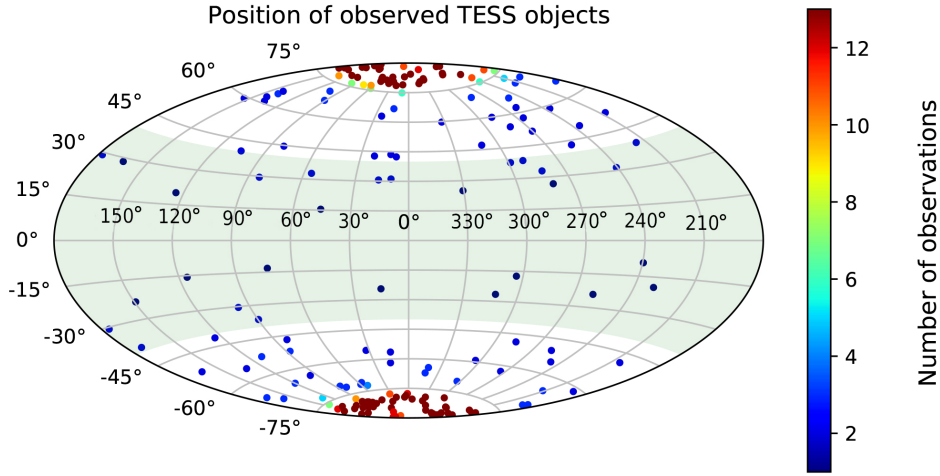


Figure 3.2: Position of each observed objects in the ecliptic frame, colour-coded to show the number of times the object is observed by TESS. The green region corresponds to the part of the sky that CHEOPS will be able to observe.

## 3.2 Simulations of Ofir systems

Figure 3.3 shows the distribution of TTV amplitudes of the objects from the Ofir catalogue. In this histogram, all systems with amplitude lower than the error are filtered out as they would be undetectable and thus not of interest in this paper. It is clear that many systems show a small TTV signal but there are a significant portion of systems which show a higher amplitude.

## 3.3 TTV signals from TESS objects

Figure 3.4 shows an example of a system observed over a short duration where there is a curvature but due to the short timescale the TTV amplitude is smaller than the error. This system may show a TTV signal if observed for a longer time. The same figure shows a TTV signal for a system at one of the ecliptic poles where the variation is clearly visible. The zero level corresponds to the average transit time in order to see the variations more clearly. The distribution of TTV amplitudes are found in figure 3.3 where amplitudes below the error are filtered out as they would

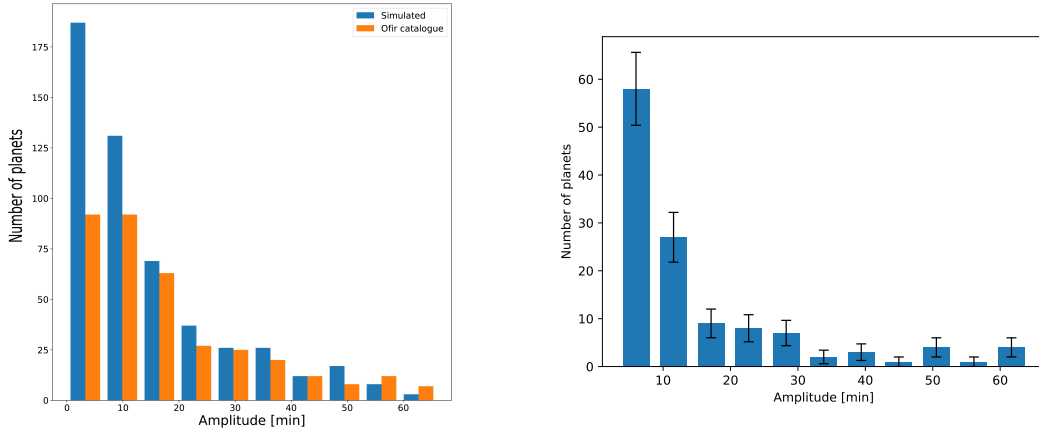


Figure 3.3: Left panel: Histogram of the amplitudes of the simulated Kepler systems selected by Ofir et al. and the amplitudes of these systems from the results of Ofir et al. Right panel: Histogram of the amplitudes of the simulated TESS objects

not be detectable and thus not of interest in the paper. It is easy to see that low or no TTV signals dominate but there are some planets showing substantial TTV signals. The location on the sky of these objects are shown in figure 3.5. Many planets at the ecliptic poles show some form of TTV signal but few planets outside do. This is due to the short observation period outside the poles. In the range that CHEOPS are able to observe 4 out of 52 planets showing a TESS detectable TTV signal. At the ecliptic poles, 30 out of 105 planets show a TESS detectable TTV signal. Over the whole sky 121 out of 474 planets in 189 systems show a TESS detectable TTV signal.

### 3.4 Error analysis

Figure 3.6 shows the amplitude of the systems as a function of the error of the amplitude. Also shown is a line where  $y=x$  where the amplitude is equal to the error. The systems with an amplitude lower than the error cannot be considered as the signal can be due to noise.

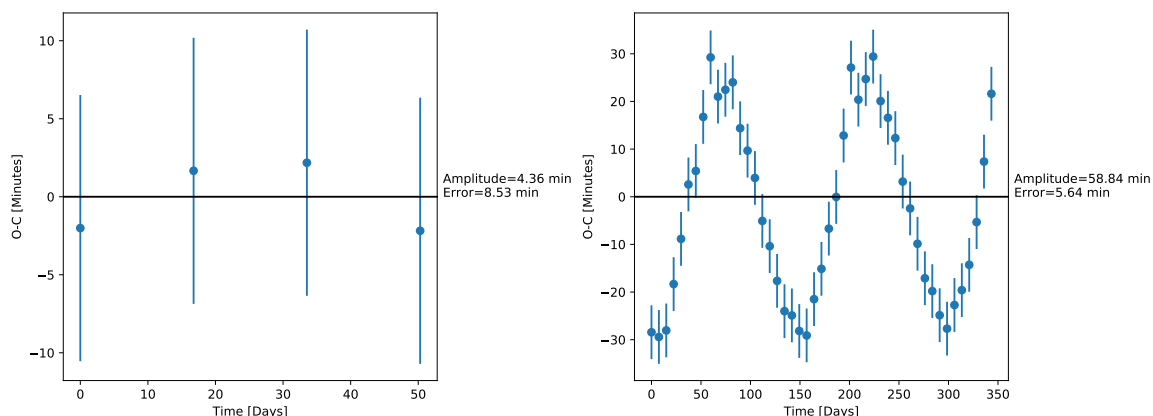


Figure 3.4: Observed-Calculated transit time as a function of time. The zero level corresponds to average transit time. The left panel shows a system which is observed for a short time which may show a TTV signal if observed over a longer time. The right panel shows a system located at the poles where the TTV signal are clearly seen.

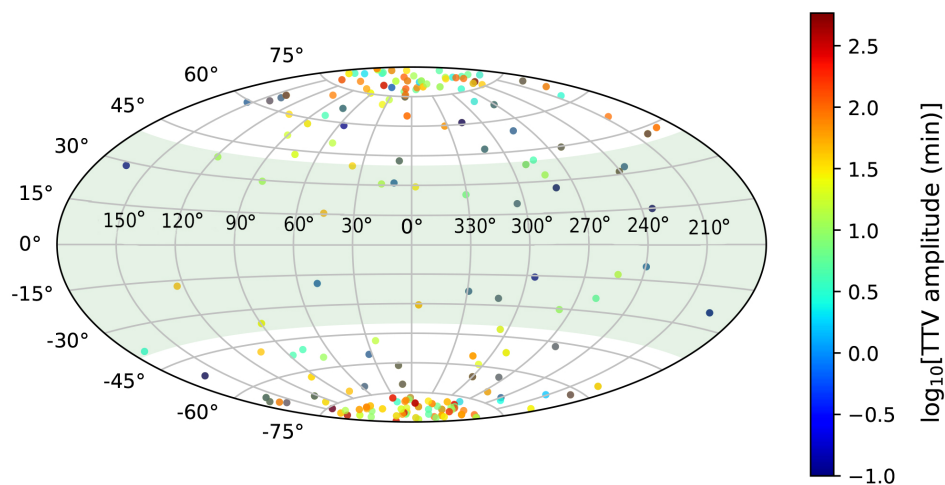


Figure 3.5: Position in the sky in the ecliptic frame and the TTV amplitude.

### 3.5 Stability simulations

Figure 3.7 shows two examples of the results from the stability simulations. For all stability tested systems most have not shown any instability within  $10^5/10^6$  years.

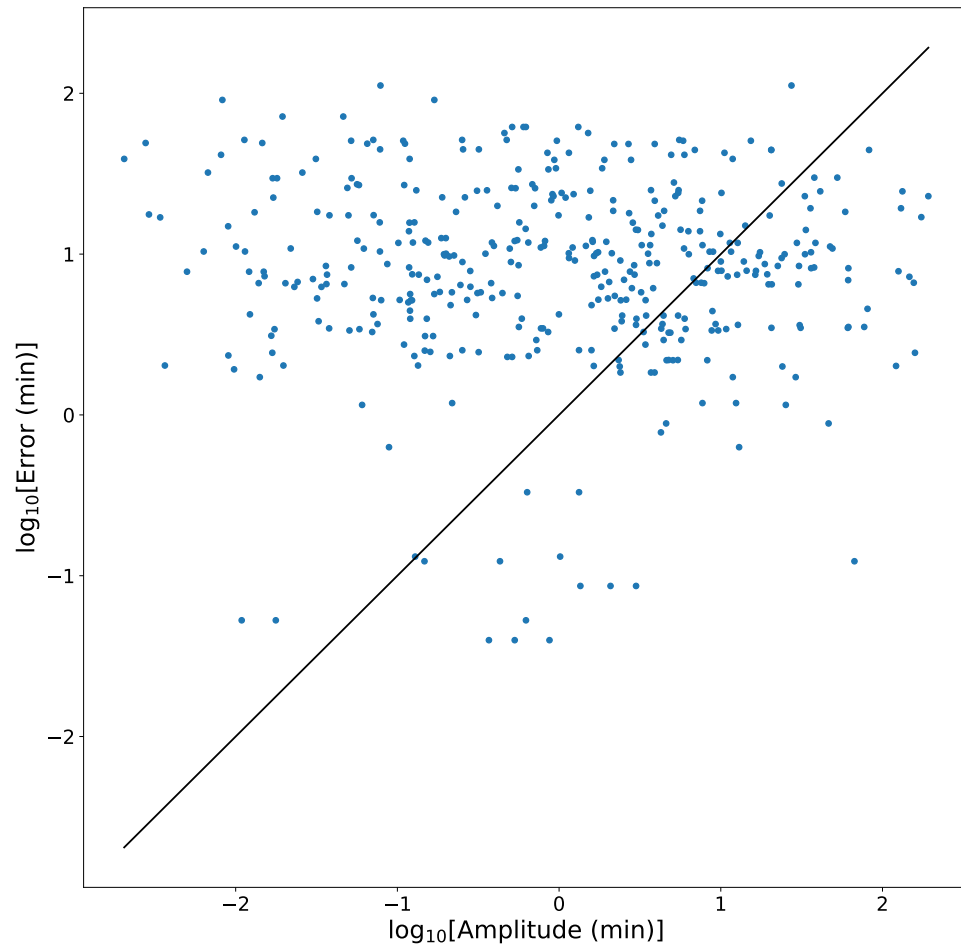


Figure 3.6: Amplitude of the TTV signal plotted against the error where  $y=x$  is marked with a black line.

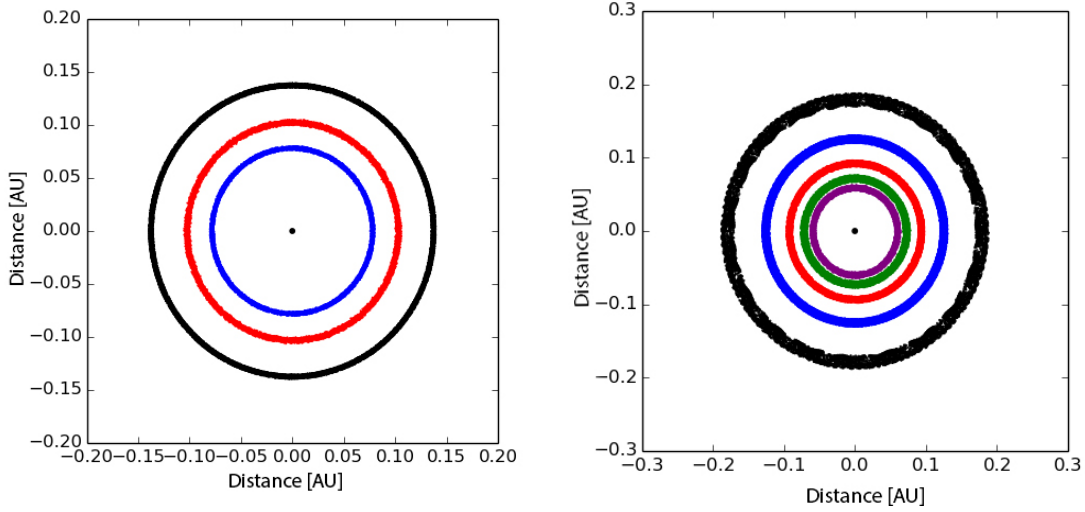


Figure 3.7: Stability simulations for two systems over  $10^6$  years. Left panel: Orbits of the planets in a three planet system. Right panel: Orbits of the planets in a five planet system

The systems were chosen from the separation obtained from the Hill radius where a lower separation means a higher chance of instability. Each system is simulated for 10 million orbits of the planet with the highest period. This creates a situation where most systems are simulated for  $10^5$  to  $10^6$  years. A few systems did eject a planet which probably is the result of a close encounter. These systems were then re-simulated using IAS15 which showed that they were in fact stable and the ejection was a result from WHFast not being able to handle the close encounters. It can be noted that all of the planets are very close to the host star. This is a result from Kepler’s third law which states that the orbital period squared is proportional to the radius of the orbit cubed, thus a low periods give low radii.

### 3.6 Amplitude dependence on time

A system was selected where the TTV signal was as close to a sine curve as possible. This curve can be seen in the left panel of figure 3.8. The resulting amplitude curve with the corresponding analytical amplitude can be seen in the right panel in figure 3.8. The TTV curve is close to a sine curve but it is slightly shifted to the right, despite this, both curves show the same behavior. Initially the curve decreases with



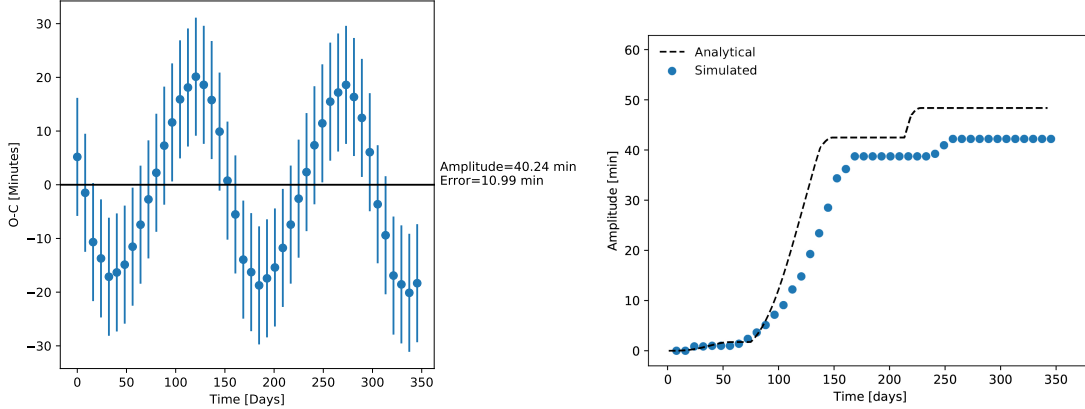


Figure 3.8: Left panel: TTV curve of selected system for amplitude analysis. Right panel: Amplitude curve of selected system where the simulated amplitude are represented by blue dots and the analytical amplitude by the dotted line.

very low curvature. After the peak the curve increases again and after it goes past the initial point the amplitude increases greatly until it plateaus. As the TTV curve is not a perfect sinusoidal curve the analytical amplitude is correct for low time but increases too quickly and overshoots the simulated value.

As the TTV amplitude increases over time a system not showing TTVs on a short timescale might show TTVs on a longer timescale such as when observed by CHEOPS. This gives an estimation of which systems are interesting and can be selected for follow up observations.

### 3.7 Objects in CHEOPS observation range

In the left panel in figure 3.9 the TTV amplitude of one planet in a system is correlated to the radius of the adjacent planet in the system. This is to see a relation between TTV amplitude of one planet and radius/mass of another as the TTV amplitude is dependent on the mass of the other planet. It is shown that the amplitude increases for systems where the period ratio is close to 1 and decreases for higher ratios. Also in general the amplitude is lower for lower radius with the highest at about  $2.5 R_{\oplus}$  which might be that there are a large amount of planets in this radius range which increases the chances of seeing a high TTV amplitude. In the right panel the TTV amplitude increases for low period ratio and decreases for higher. At around period ratio = 2 there are a visible peak in amplitude which originates from

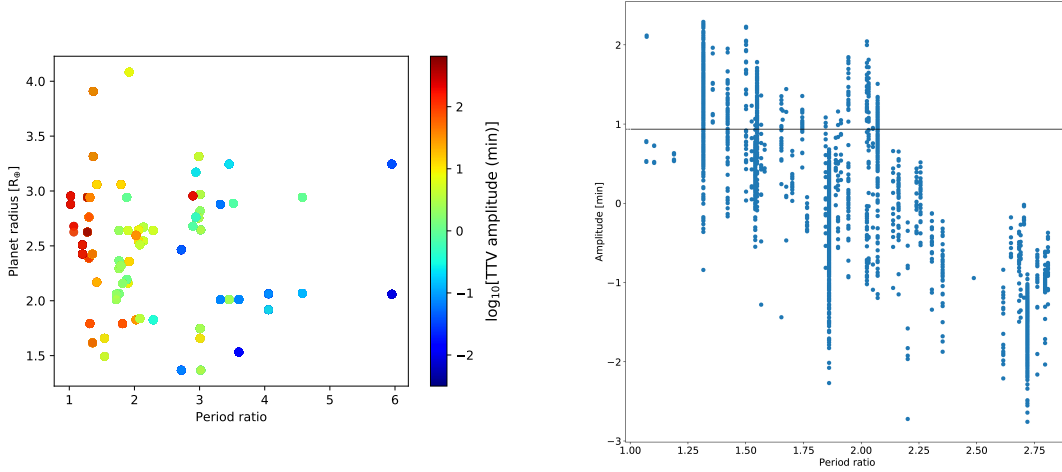


Figure 3.9: Left panel: Planet radius plotted against period ratio where the amplitude of the CHEOPS-detected TTV signal is color coded. Right panel: Amplitude of the CHEOPS-detected TTV signal plotted against period ratio where the median TTV amplitude error is represented as a black horizontal line.

the 2:1 resonance. In total 3000 systems were created which contains 6800 planets. Of these, 6798 planets showed some form of TTV signal and out of these 3291 planets showed a CHEOPS detectable TTV signal.

In figure 3.10 the probability to find a detectable TTV signal during 1 year of observations based on the planet radius and period ratio is shown. The reason for the duration of 1 year is that this is the maximum time TESS will spend looking at one spot in the sky, the ecliptic poles. As expected low period ratio gives a higher probability and higher period ratio show very low probability to get a detectable signal. From these results, if a planet with, for example, radius around 2.5  $R_{\oplus}$  and a period ratio of 3 is found there are a large probability that that planet will show a detectable TTV signal.

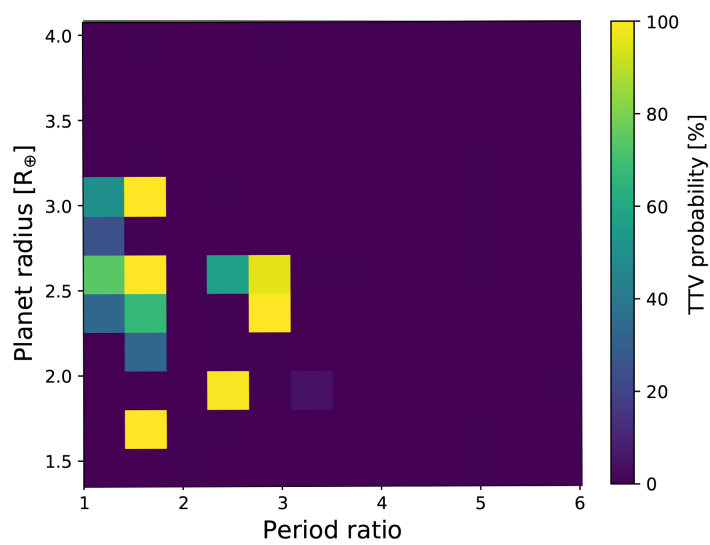


Figure 3.10: 2d histogram of CHEOPS-detectable TTV probability depending on planet radius and period ratio.

# Chapter 4

## Discussion

These results may be used to give an approximated upper limit of how many systems in a given sample will show TTV signals. The systems can be observed in more detail by, for example, the CHaracterising ExOPlanets Satellite, CHEOPS, or the James Webb Space Telescope, JWST. CHEOPS is not able to measure stars at the poles and is limited to the sky around the plane of Earth's orbit which is marked as a green area in figure 3.2 and 3.5. As TESS will have coverage of the poles for the whole year many of the systems showing TTV signals will be located at the poles. The JWST on the other hand will be able to look at the poles. CHEOPS will be used to further study already found systems in order to more accurately measure the radii of the exoplanets. It can be used to obtain more transits and thus give more opportunities to measure TTVs. JWST will study the spectrum of a star and compare it to the spectrum of the same star when an exoplanet is transiting. With high enough precision this will make it possible to approximate the composition of an potential atmosphere around the exoplanet. To determine the atmosphere the surface gravity of the planet needs to be known which comes from the mass and radius of the planet. The radius can be obtained from the transit method while the mass may be obtained from a TTV signal or the radial velocity method.

The way of determining the amplitude of the TTVs are very simplistic. It is simply an average of the highest and lowest value. This can be improved by some kind of fit as the TTV signals are often in the shape of a sine curve but due to time constraints this paper does not include this way of determining the amplitude.

Even though these results show that CHEOPS might be able to find a planet showing TTV signals in every other system the real number is probably lower. An approximation used in this paper is that the inclination is  $90^\circ$  for all planets. This creates systems where all planets are transiting which is not the case for real systems.

Also for some TTV signals the super-period, the period of the TTV signal, is longer than the observational baseline. In these cases fitting a sinusoid is meaningless as the signal won't have time to show the full period.

The systems are checked for stability to ensure that they are realistic with WHFast, IAS15 and by comparing the distribution of amplitudes from the TESS objects with the objects from the Ofir et al. (2018) paper. Using WHFast, most systems tested for stability show no sign of instability within  $10^5/10^6$  years. The few that do are simulated using IAS15 where they are shown to be stable and the instability from WHFast can be considered an effect of WHFast's inability to handle close encounters. The artificial systems created in this paper contain some planets which show strong TTV signals, mainly positioned in the ecliptic poles as the systems positioned there are observed for a whole year which gives many more opportunities to measure the transits. The shape of the histogram in figure 3.3 are the same as in the Ofir paper. This shows that the methods of simulating systems used in this paper are viable and give reasonable results. The results from this paper does however show more low amplitude TTVs than Ofir et al. (2018) based on this the results may be a bit pessimistic but as the curve is the same for higher amplitude the results still seem reasonable.

The error calculations seem to give reasonable results (Becker et al. 2015). As seen in figure 3.6 many systems shows a TTV amplitude higher than the error but many systems shows a TTV amplitude lower than the error, for these systems no conclusion can be made as the signal could be true but it might as well be noise. The error is solely based on photon statistics and does not take things like cadence, how often the luminosity is measured or sample rate, into consideration which would otherwise increase the errors.

The error calculation done in this paper is based on the photometry of TESS. As CHEOPS will have a larger camera than TESS the error will be of a factor 2-3 lower and thus increase the chance of a detectable TTV signal.

The TTV amplitudes of objects within the CHEOPS observation range are in general low due to the short observation time of these objects. By considering figure 3.8 where the amplitude is shown to increase over time and figure 3.10 where high TTV probability can be seen for systems with low period ratio and relatively high radius/mass the long term TTV amplitude may be predicted. A transiting planet with a radius of about  $2.5 R_{\oplus}$  and a period ratio of 2 should, based on the simulations in this paper, have a fairly high probability to show a detectable TTV signal. On the other hand, for a period fraction greater than 3 the probability to find a detectable TTV signal is very low.

# Chapter 5

## Conclusions

This paper simulated TESS objects using data from the Kepler dataset and results from the paper by Sullivan et al. (2015) on a search to predict TTV signals caused by multiple planets in the systems. The results can be seen below:

- Most of the systems showing TTV signals are positioned at the ecliptic poles. This comes from the fact that TESS has continuous coverage of the poles during the year which increases the chance to see TTVs. The amplitude grows non-linearly with time and thus a system might show very low amplitude on a short timescale but very high if observed for longer.
- In the range where CHEOPS is able to observe only 4 out of 52 planets showed a TTV amplitude greater than the TESS amplitude error when simulated on a short timescale. Upon further analysis, 3291 out of 6800 planets showed a CHEOPS detectable TTV signal when simulated for 1 year. At the ecliptic poles where the JWST continuous viewing zones are located, 30 out of 105 planets showed a TESS detectable TTV signal. Over the whole sky 121 planets out of a total 474 showed a TESS detectable TTV signal.
- Based on TESS observations objects around the ecliptic, where CHEOPS will observe, will show a low TTV signal due to the low observation time. Based on further simulations these objects may show a detectable TTV signal over a longer observation period which may be predicted based on the planetary radius and period fraction of the planets in the system. Planets with a planet fraction of 2 or below and a radius of about  $2 R_{\oplus}$  or higher are likely to show a CHEOPS-detectable TTV signal. If the period fraction is above 3 the probability to see a CHEOPS-detectable TTV signal is very low.

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Simulations in this paper made use of the REBOUND code which can be downloaded freely at <http://github.com/hannorein/rebound>.

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